Automatic Shape Recognition of Type III Radio Bursts in Solar Wind Dynamical Radio Spectra

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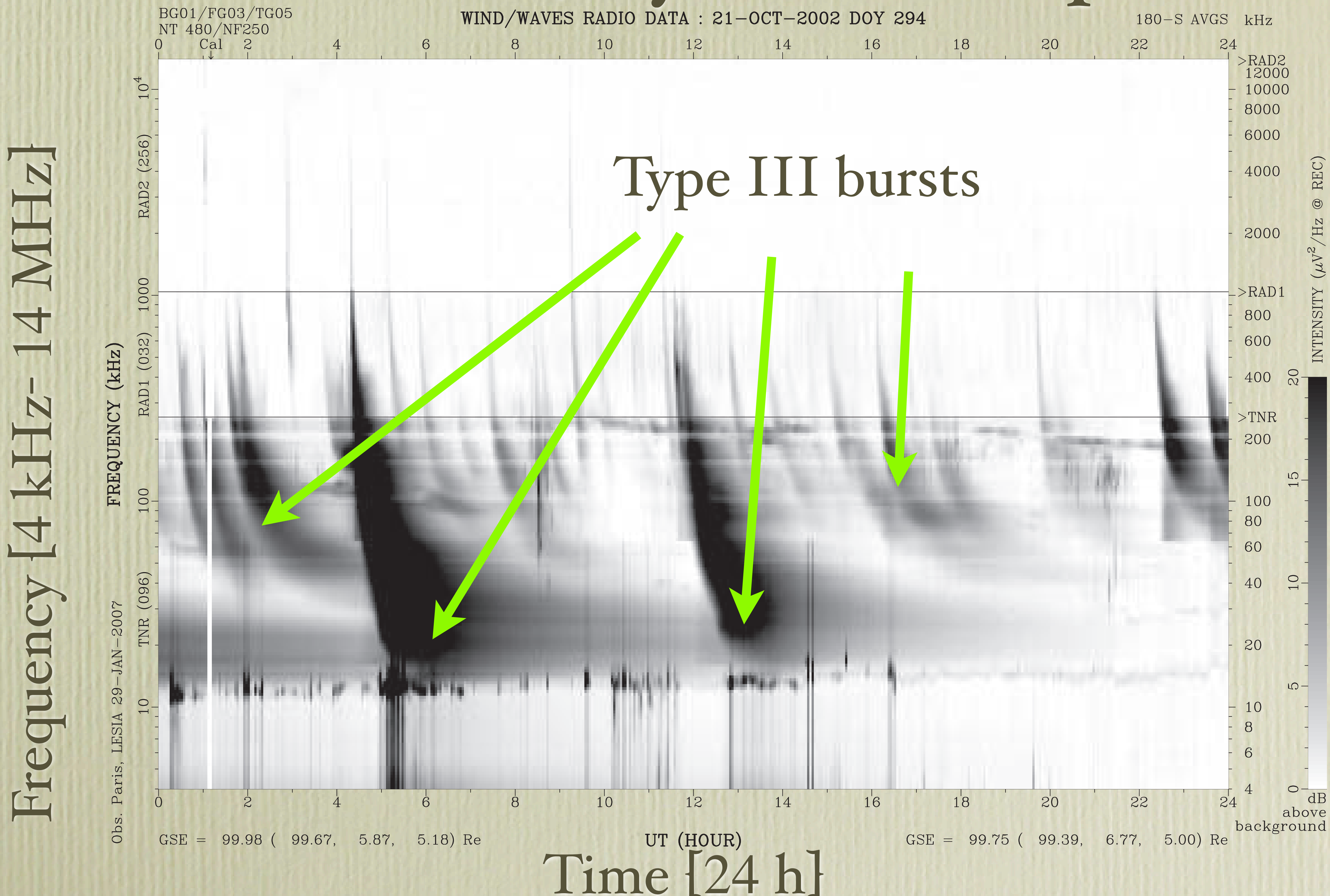
Motivation

- Large amounts of data are already recorded and stored - they continue to grow every day.
- People have no time to analyze this data - human attention has become the precious resource.
- So, we must find ways to automatically analyze the data, to automatically classify it, summarize it, to discover and characterize trends in it, to automatically flag anomalies etc.

Observations, dynamical spectrum



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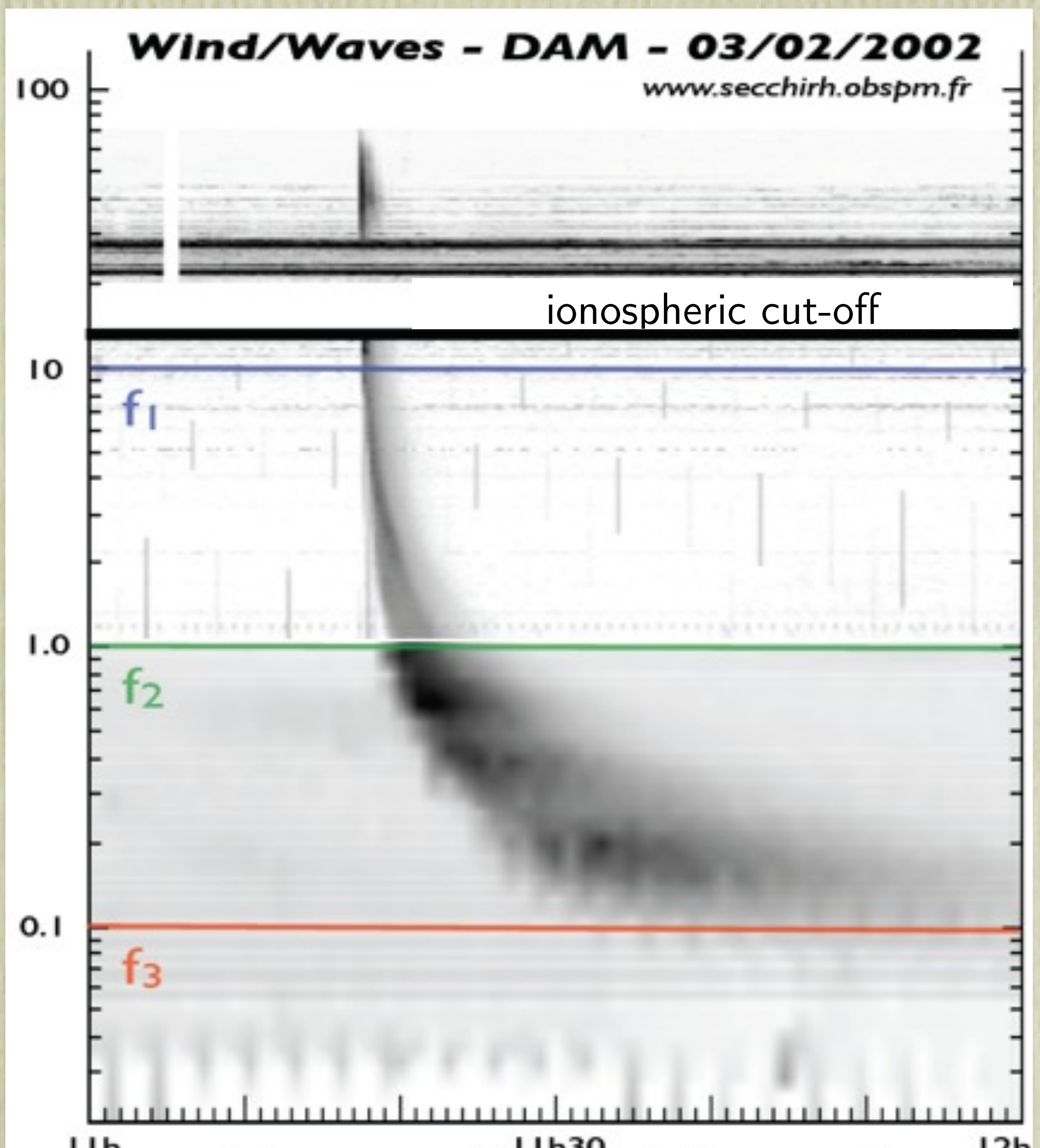
Type III Bursts



- Short (sec → hrs) & very intense (→10-14 Wm-2Hz-1) radio emissions;
- Emission frequencies decrease rapidly (GHz → kHz);
- Emission at fundamental fp or at harmonic of fp;
- Often associated with solar flares;
- Associated with the propagation of electrons supra-thermal (c/10 → c/3);

TIII's Frequency Drift*) (1/2)

- The frequency, related to the local plasma frequency ($f_{pl} \propto \sqrt{n_e}$, $n_e \propto 1/R^2$, $f_{pl} \propto 1/R$), drifts downward as the emission region rapidly propagates outward.
- Since the radio burst is generated by local plasma emission processes, radio emissions at high frequencies (high plasma densities) occur very near the Sun ~ $2R_\odot$ for 16 MHz, while those at low frequencies (low plasma densities) occur far from the Sun (~ 1 AU) for 20 kHz.
- Type III radio bursts are therefore characterized by a rapid drift to lower frequencies due to the near-relativistic speeds of the burst electrons.

*) Vidojevic S., Maksimovic M.: *Preliminary Analysis of Type III Radio Bursts from STEREO/SWAVES Data*, XV National Conference of Astronomers of Serbia, 2–5 October 2008, Belgrade, Serbia, Publ. Astron. Obs. Belgrade No. **86** (2009), 287 - 291. http://publications.aob.rs/86/pdf/287-291.pdf

TIII's Frequency Drift*) (2/2)

- For about 100 bursts we have computed the frequency drift rates obtained from all the maxima of the power spectral density profiles at each of the covered frequencies. The profiles are fitted by Gram-Charlier type A function.

- Obtained maxima are further aproximated by linear function in log-log scale.

- **$df/dt = -10^{a} f^{\alpha}$**. The negative sign denotes that the starting frequency is observed to drift from high to low values. The least square fit of a straight line through all of observed maxima gives:

- **$\alpha = 1.80 \pm 0.05$ and $a = -1.70 \pm 0.03$.**

Spectra Shape Modelling*)

- The choice of the best-suited statistical distribution for data modelling is not a trivial issue;
- Unless a sound theoretical background exists for selecting a particular distribution, one will usually try to test various candidates and select a distribution based on its fit to the observed data;
- It is more efficient to define a sufficiently general family that can be used for this purpose.

*) S. Vidojevic Shape Modelling with Family of Pearson Distributions, 9th SerbianConference on Spectral Line Shapes in Astrophysics, Banja Koviljaca, Serbia, May 13-17, 2013, Book of abstracts, p. 52, http://www.scslsa.matf.bg.ac.rs/Book_of_abstracts_9thSCSLSA.pdf

Pearson system - great diversity of shapes:

- unimodal, bimodal, U-shaped, J-shaped and monotone probability distribution functions,
- ...which may be symmetric and asymmetric, concave and convex,
- ...with smooth, abrupt, truncated, long, medium or short tails.

Pearson system*)

- First derivative of probability density function:

$$\frac{1}{f(x)}\frac{\mathrm{d}f(x)}{\mathrm{d}x} = -\frac{a+x}{c_0 + c_1 x + c_2 x^2}$$

- Asymmetry (As^2= β_1)

$$\beta_1 = \frac{\mu_3^2}{\mu_2^2}$$

- Excess (β_2)

$$\beta_2 = \frac{\mu_4}{\mu_2^2}$$

Using only 2 parameters: Squared Asymmetry (β_1) and Excess (β_2), calculated from observations, Type of Pearson distribution can be retrieved.

*) Pearson, K.: 1895, *Contributions to the Mathematical Theory of Evolution. II. Skew Variation in Homogeneous Material.* Philosophical Transactions of the Royal Society of London, **186**, 343 – 414

Method of moments

$$c_0 = (4\beta_2 - 3\beta_1)(10\beta_2 - 12\beta_1 - 18)^{-1}\mu_2$$

$$a = c_1 = \sqrt{\beta_1}(\beta_2 + 3)(10\beta_2 - 12\beta_1 - 18)^{-1}\sqrt{\mu_2}$$

$$c_2 = (2\beta_2 - 3\beta_1 - 6)(10\beta_2 - 12\beta_1 - 18)^{-1}$$

$$\kappa = \frac{1}{4}c_1^2(c_0 c_2)^{-1} = \frac{1}{4}\beta_1(\beta_2 + 3)^2(4\beta_2 - 3\beta_1)^{-1}(2\beta_1 - 6)^{-1}$$

Classification

I: $\kappa < 0$

II: $\beta_1 = 0,\ \beta_2 < 3$

III: $2\beta_2 - 3\beta_1 - 6 = 0$

IV: $0 < \kappa < 1$

V: $\kappa = 1$

VI: $\kappa > 1$

VII: $\beta_1 = 0,\ \beta_2 > 3$

Method of Maximum Likelihood

Likelihood function

$$L(\boldsymbol{\theta}|\mathbf{x}) \equiv f(\mathbf{x}|\boldsymbol{\theta}) = \prod_{i=1}^{n} f_i(x_i|\boldsymbol{\theta})$$

applying logarithm, one obtain:

$$\mathcal{L}(\boldsymbol{\theta}|\mathbf{x}) = \ln L(\boldsymbol{\theta}|\mathbf{x}) = \sum_{i=1}^{n} \ln f_i(x_i|\boldsymbol{\theta})$$

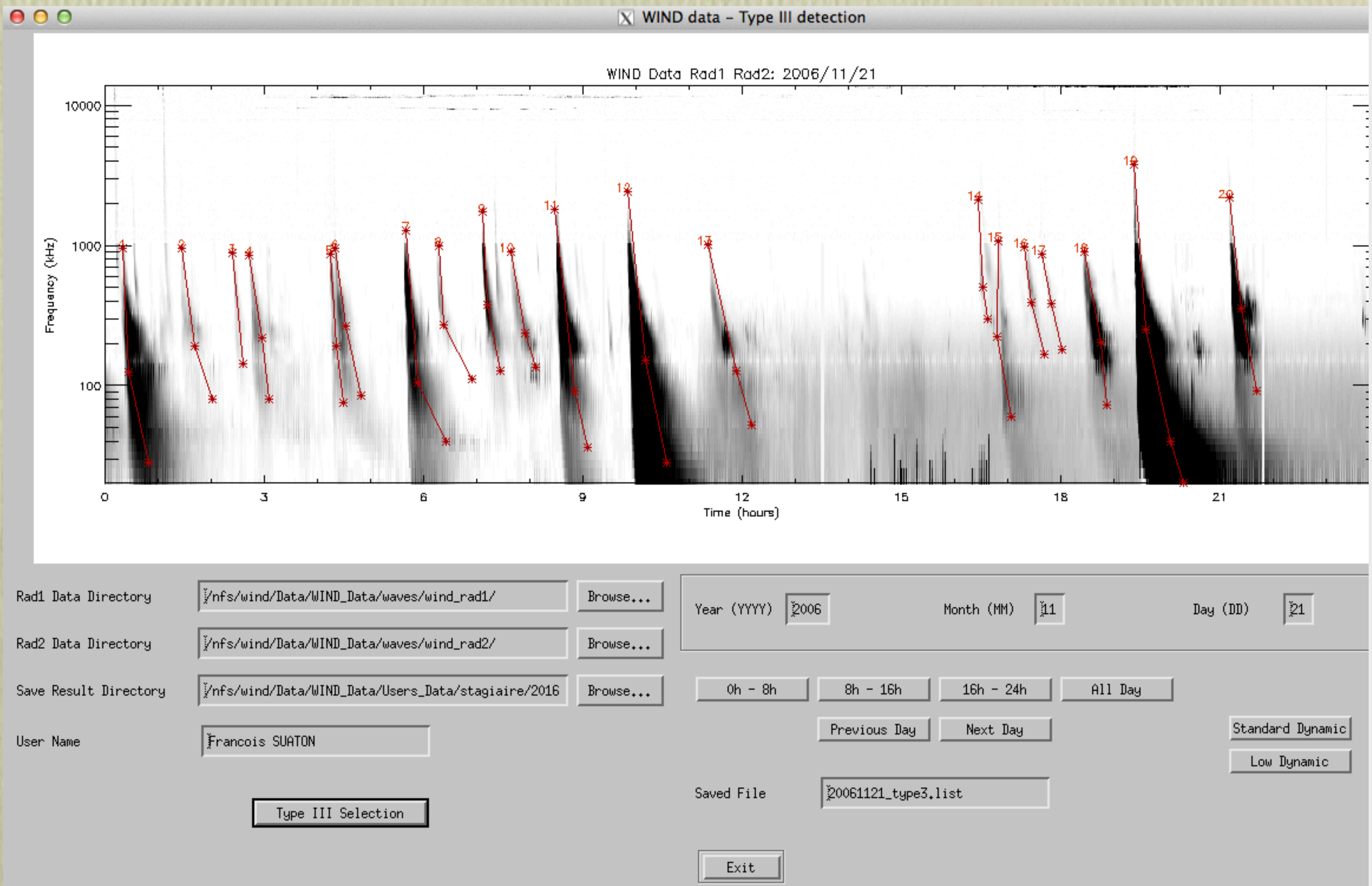
Looking for $\boldsymbol{\theta}^*$

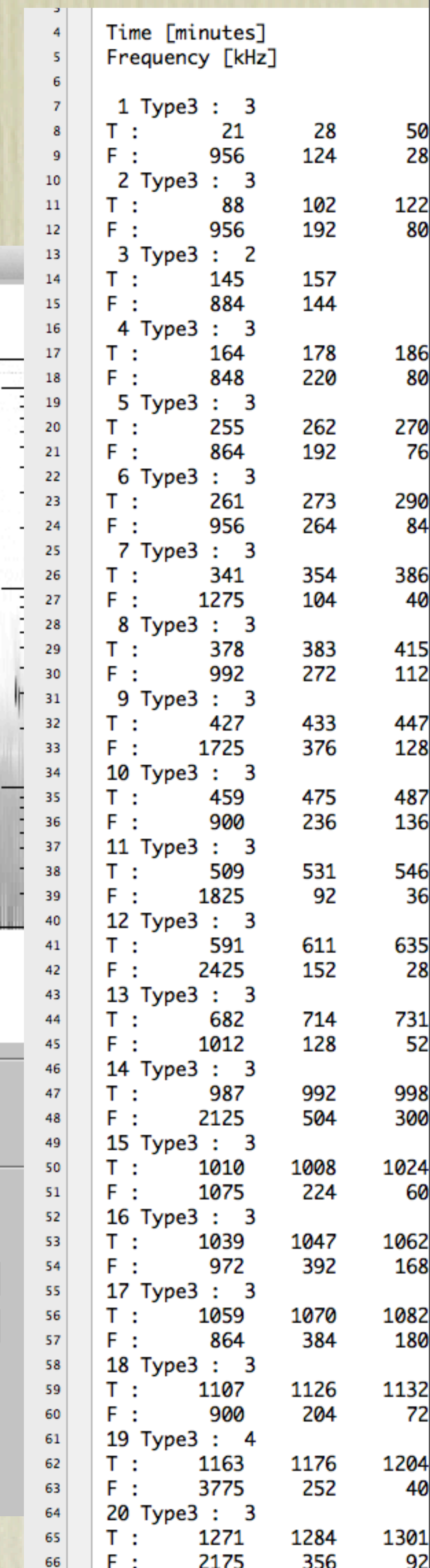
- Looking for $\boldsymbol{\theta}^*$ which maximizes likelihood

$$\mathcal{L}(\boldsymbol{\theta}^*|\mathbf{x}) = \max_{\boldsymbol{\theta}} \mathcal{L}(\boldsymbol{\theta}|\mathbf{x}) = \max_{\boldsymbol{\theta}} \sum_{i=1}^{n} \ln f_i(x_i|\boldsymbol{\theta})$$

- It is not possible to solve this task analytically, thus, we apply numerical methods of optimization.

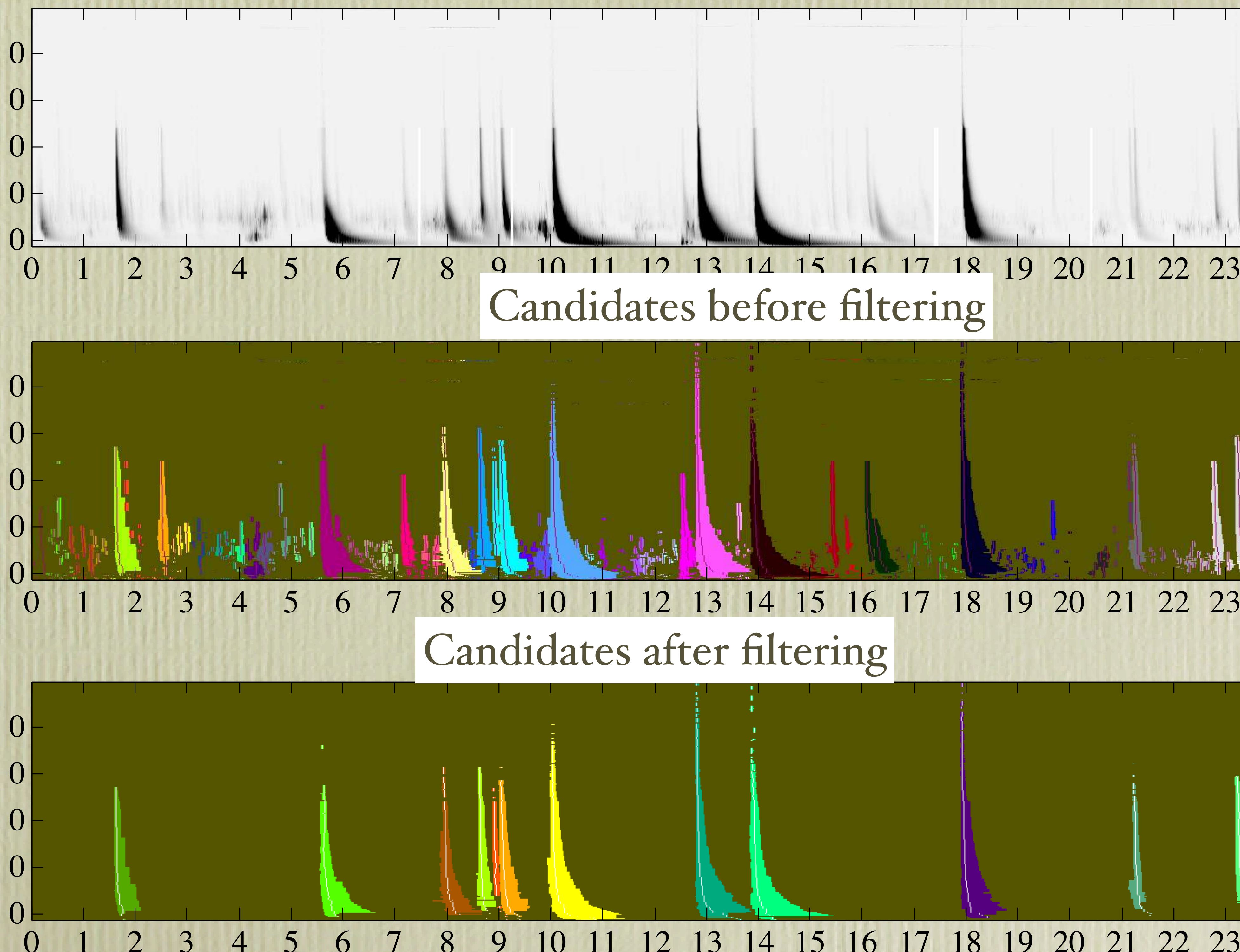
Manual Detection





Example 1

Original data 19971123



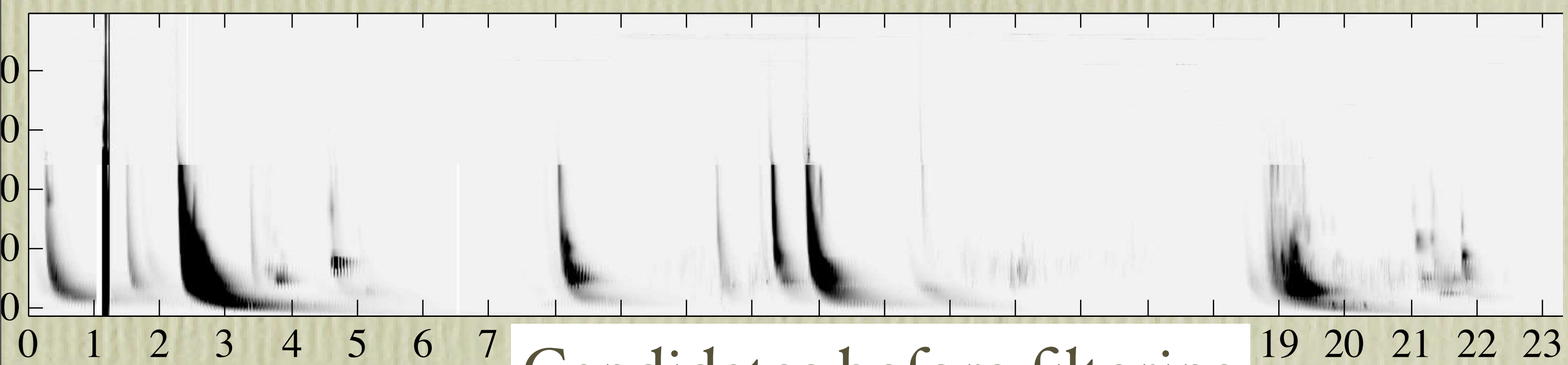
Processing date: 19971123 with parameters:

find_peaks_cutoff: 1.8
find_peaks_slope: 0.2
find_peaks_peakchkdist: 5
find_hills_maxdistforcontnext: 3
find_hills_overlaptol: -4
find_hills_maxdistforcontall: 15
find_hills_peakvalchangetol: [0.5 2]
filter_hills_noisedetect: [4 100 10]
filter_hills_maxdistforcont: 20
filter_hills_minhilllen: 50
filter_hills_notbelowfreq: 1.1
filter_hills_shapecheck: [1 1]
filter_hills_delreport: [0 0 0 0 0 1 1]
image_hills_peakvalue: 50

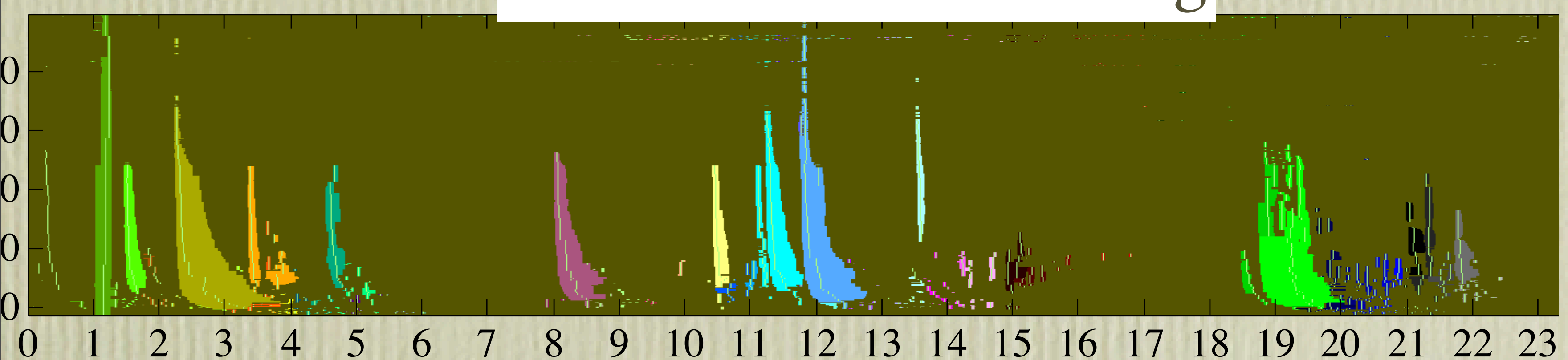
Hills before filtering: 194
Hills after filtering: 12
Results saved to data/19971123_res.mat

Example 2

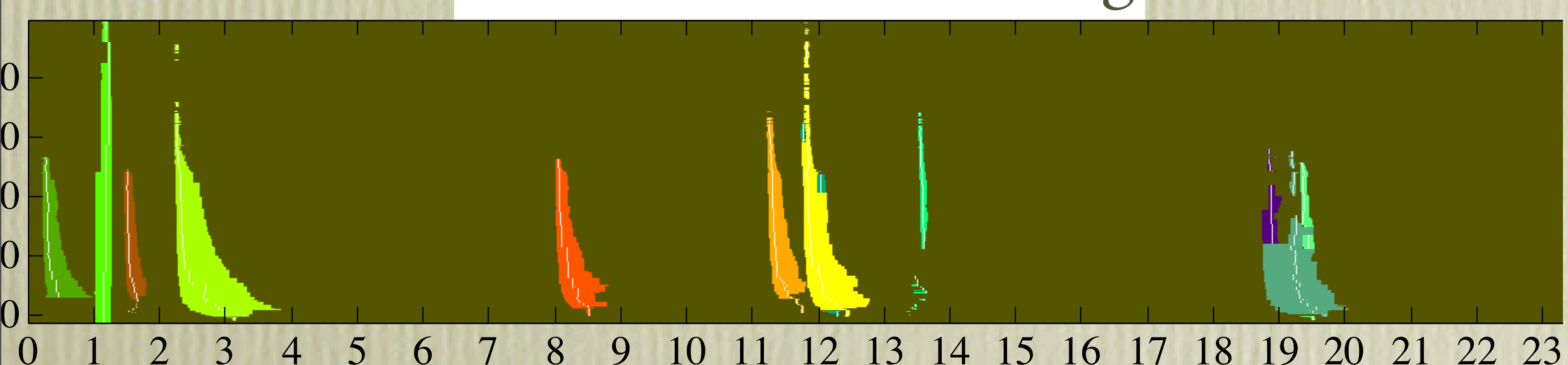
Original data 20020703



Candidates before filtering



Candidates after filtering



- Processing date: 20020703 with parameters:

find_peaks_cutoff: 1.8
find_peaks_slope: 0.2
find_peaks_peakchkdist: 5
find_hills_maxdistforcontnext: 3
find_hills_overlaptol: -4
find_hills_maxdistforcontall: 15
find_hills_peakvalchangetol: [0.5 2]
filter_hills_noisedetect: [4 100 10]
filter_hills_maxdistforcont: 20
filter_hills_minhilllen: 50
filter_hills_notbelowfreq: 1.1
filter_hills_shapecheck: [1 1]
filter_hills_delreport: [0 0 0 0 0 1 1]
image_hills_peakvalue: 50

Hills before filtering: 164
Take out hill 126 at 19.08 h : not convex (-3.719713)
Hills after filtering: 12
Results saved to data/20020703_res.mat

More data - satellites:

- WIND spacecraft, launched 1994, still operating.
- STEREO A and B, launched 2006, still operating.
- Solar Probe Plus, to be launched in 2018.
- Solar Orbiter, to be launched in 2019.

Literature

- Vidojevic S., Maksimovic M.: P*reliminary Analysis of Type III Radio Bursts from STEREO/SWAVES Data*, XV National Conference of Astronomers of Serbia, 2–5 October 2008, Belgrade, Serbia, Publ. Astron. Obs. Belgrade No. **86** (2009), 287 - 291. http://publications.aob.rs/86/pdf/287-291.pdf
- S. Vidojevic Shape Modelling with Family of Pearson Distributions, 9th SerbianConference on Spectral Line Shapes in Astrophysics, Banja Koviljaca, Serbia, May 13-17, 2013, Book of abstracts, p. 52, http://www.scslsa.matf.bg.ac.rs/Book_of_abstracts_9thSCSLSA.pdf
- Pearson, K.: 1895, *Contributions to the Mathematical Theory of Evolution. II. Skew Variation in Homogeneous Material.* Philosophical Transactions of the Royal Society of London, **186**, 343 – 414.
- Sir Ronald Aylmer Fisher (1890-1962) for the first time presented the idea in 1912 (when he was 22 years old) in the article: On an absolute criterion for fitting frequency curves, Messenger of Mathematics (1912), **41,** 155-160.

Thursday, August 24, 17